

2004

# Evidence for Correlated Titanium and Deuterium Depletion in the Galactic ISM

Jason X. Prochaska

*University of California - Santa Cruz*

Todd M. Tripp

*University of Massachusetts - Amherst*

Christopher Howk

*University of California - San Diego*

Follow this and additional works at: [http://scholarworks.umass.edu/astro\\_faculty\\_pubs](http://scholarworks.umass.edu/astro_faculty_pubs)



Part of the [Astrophysics and Astronomy Commons](#)

---

Prochaska, Jason X.; Tripp, Todd M.; and Howk, Christopher, "Evidence for Correlated Titanium and Deuterium Depletion in the Galactic ISM" (2004). *Astronomy Department Faculty Publication Series*. Paper 30.

[http://scholarworks.umass.edu/astro\\_faculty\\_pubs/30](http://scholarworks.umass.edu/astro_faculty_pubs/30)

This Article is brought to you for free and open access by the Astronomy at ScholarWorks@UMass Amherst. It has been accepted for inclusion in Astronomy Department Faculty Publication Series by an authorized administrator of ScholarWorks@UMass Amherst. For more information, please contact [scholarworks@library.umass.edu](mailto:scholarworks@library.umass.edu).

## EVIDENCE FOR CORRELATED TITANIUM AND DEUTERIUM DEPLETION IN THE GALACTIC ISM

JASON X. PROCHASKA<sup>1,2</sup>, TODD M. TRIPP<sup>3</sup>, J. CHRISTOPHER HOWK<sup>2,4</sup>  
*Accepted to the Astrophysical Journal Letters: December 28, 2004*

### ABSTRACT

Current measurements indicate that the deuterium abundance in diffuse interstellar gas varies spatially by a factor of  $\sim 4$  among sightlines extending beyond the Local Bubble. One plausible explanation for the scatter is the variable depletion of D onto dust grains. To test this scenario, we have obtained high signal-to-noise, high resolution profiles of the refractory ion Ti II along seven Galactic sightlines with D/H ranging from 0.65 to  $2.1 \times 10^{-5}$ . These measurements, acquired with the recently upgraded Keck/HIRES spectrometer, indicate a correlation between Ti/H and D/H at the  $> 95\%$  c.l. Therefore, our observations support the interpretation that D/H scatter is associated with differential depletion. We note, however, that Ti/H values taken from the literature do not uniformly show the correlation. Finally, we identify significant component-to-component variations in the depletion levels among individual sightlines and discuss complications arising from this behavior.

*Subject headings:* cosmology: observations, ISM: evolution, ISM: abundances, galaxy: abundances

### 1. INTRODUCTION

In standard big bang nucleosynthesis (BBN), deuterium is the most sensitive baryometer of the light elements (Schramm & Turner 1998). Measurements of its primordial value in quasar absorption line systems have placed tight constraints on the baryonic mass density  $\Omega_b$  (Burles & Tytler 1998; O'Meara et al. 2001; Kirkman et al. 2003). The inferred  $\Omega_b$  value is in impressive agreement with the value derived from CMB measurements of the WMAP (e.g. Spergel et al. 2003) and other microwave experiments lending confidence to both the deuterium measurements and BBN theory. In addition to its cosmological significance, deuterium is an important tracer of chemical evolution. Deuterium is astrated within stellar cores, and there are no known means of producing significant amounts of D other than BBN. The evolution of D/H, therefore, tracks the global history of the gas astration within a galaxy (e.g. Chiappini, Renda, & Matteucci 2002).

It is in this context that measurements of D/H within the Milky Way have impact: (1) the upper bound to the Galactic D/H value sets a lower limit to the primordial D/H value; and (2) a comparison of Galactic D/H with the primordial D/H value describes the chemical evolution history of the Galaxy. Since the first measurements of Galactic D/H with *Copernicus* (Rogerson & York 1973), surveys for Galactic D/H have been pursued on each succeeding UV observatory bearing a high resolution spectrometer (e.g. Linsky et al. 1995; Jenkins et al. 1999; Hoopes et al. 2003; Wood et al. 2004). The most remarkable conclusion of these efforts is that beyond the Local Bubble the D/H values have significant scatter; D/H ranges from 5 to  $20 \times 10^{-6}$ . This dispersion is larger than the estimated statistical error and is unlikely to be systematic error as-

sociated with the data analysis (see references in Table 1 for detailed and thorough analyses of the measurement uncertainties). At present, it seems that the D/H ratio has intrinsic scatter within the Milky Way.

Recently, observations with the Interstellar Medium Absorption Profile Spectrograph (IMAPS) and the *FUSE Observatory* have extended D/H measurements to large distances from the Sun and correspondingly higher  $N(\text{HI})$  values. Although only a handful of measurements exist to date, Wood et al. (2004) and others have noted that the few sightlines that probe the greatest distances ( $> 500\text{pc}$ ) have a central value and dispersion that are significantly smaller than sightlines at intermediate distances ( $20\text{pc} < d < 500\text{pc}$ ). One possible interpretation of these trends is that D is significantly depleted onto dust grains (e.g. Jura 1982). Draine (2004a,b) has examined the principal mechanisms of D adsorption and grain destruction and argues that it is at least plausible that D would be preferentially depleted. He proposed testing the depletion hypothesis by comparing the D/H values with abundances of refractory elements like Fe, Ni, and Ti (i.e. species that are highly prone to depletion onto grains) along the same sightlines.

We have initiated a program to obtain high signal-to-noise (S/N), high resolution observations of Ti II profiles for sightlines with accurate Galactic D/H measurements. Previous surveys of Ti have demonstrated that it is highly refractory, presumably because of its large condensation temperature (e.g. Stokes 1978; Lipman & Pettini 1995). Therefore, a measurement of Ti/H in the ISM assesses the depletion level along that sightline. Of additional importance,  $\text{Ti}^+$  is the dominant ion in H I regions. Unlike  $\text{Na}^0$  and  $\text{Ca}^+$  the ionization potential of  $\text{Ti}^+$  is  $\approx 1\text{Ryd}$  and  $\text{Ti}^+$  is predominantly shielded from ionizing photons. Furthermore,  $\text{Ti}^+$  has a charge exchange reaction rate with hydrogen that is significantly greater than many other ions (e.g.  $\text{Fe}^+$ ,  $\text{Si}^+$ ; Kingdon & Ferland 1996). We expect is less sensitive to photoionization effects and Ti II should trace the velocity profiles of H I and D I gas. Therefore, high resolution Ti II profiles are likely to be better suited than, e.g., Fe II, for constraining the fits of D I and assessing the likelihood of deuterium line saturation in high  $N(\text{HI})$  sightlines.

<sup>1</sup> UCO/Lick Observatory; University of California, Santa Cruz; Santa Cruz, CA 95064, xavier@ucolick.org

<sup>2</sup> Visiting Astronomer, W. M. Keck Telescope. The Keck Observatory is a joint facility of the University of California and the California Institute of Technology.

<sup>3</sup> Department of Astronomy; University of Massachusetts; 710 N. Pleasant St.; Amherst, MA 01003-9305

<sup>4</sup> Department of Physics, and Center for Astrophysics and Space Sciences, University of California, San Diego, C-0424, La Jolla, CA 92093-0424

TABLE 1  
SUMMARY OF OBSERVATIONS

Target	$\epsilon$ Ori	$\delta$ Ori	HD 191877	HD 195965	BD+28°4211	$\iota$ Ori	Feige 110
$B$	1.71	2.01	6.22	6.60	13.01	2.53	13.01
Obs (UT)	08Sep2004	08Sep2004	08Sep2004	08Sep2004	06Oct2004	08Sep2004	09Sep2004
$t_{exp}$ (s)	9	13	1800	1200	1800	9	1800
S/N <sup>a</sup>	370	355	350	250	145	280	120
$\log[N(\text{HI})/\text{cm}^{-2}]$	$20.40 \pm 0.08$	$20.19 \pm 0.03$	$21.05 \pm 0.05$	$20.95 \pm 0.03$	$19.85 \pm 0.02$	$20.16 \pm 0.10$	$20.14 \pm 0.07$
D/H <sup>b</sup>	$0.646 \pm 0.38$	$0.736 \pm 0.12$	$0.776 \pm 0.20$	$0.851 \pm 0.17$	$1.380 \pm 0.10$	$1.413 \pm 0.88$	$2.138 \pm 0.43$
Ref	1	2	3	3	4	1	5
$\log[N(\text{Ti II } \lambda 3230)]$	$< 11.55$	$< 11.46$	$12.26 \pm 0.03$	$11.97 \pm 0.05$	$< 11.57$	$< 11.67$	$< 11.77$
$\log[N(\text{Ti II } \lambda 3242)]$	$11.50 \pm 0.05$	$11.28 \pm 0.06$	$12.25 \pm 0.02$	$12.06 \pm 0.02$	$< 10.99$	$11.20 \pm 0.07$	$11.69 \pm 0.05$
$\log[N(\text{Ti II } \lambda 3384)]$	$11.37 \pm 0.04$	$11.09 \pm 0.06$	$12.23 \pm 0.02$	$12.00 \pm 0.02$	$11.08 \pm 0.08$	$11.34 \pm 0.03$	$11.54 \pm 0.04$
$\log[N(\text{Ti II})]^c$	$11.40 \pm 0.03$	$11.15 \pm 0.04$	$12.24 \pm 0.02$	$12.02 \pm 0.02$	$11.08 \pm 0.08$	$11.31 \pm 0.03$	$11.59 \pm 0.03$
$\log(\text{Ti}/\text{H})$	$-9.00 \pm 0.09$	$-9.05 \pm 0.05$	$-8.81 \pm 0.05$	$-8.93 \pm 0.03$	$-8.77 \pm 0.08$	$-8.85 \pm 0.10$	$-8.55 \pm 0.09$

<sup>a</sup>Empirically measured per pixel at 3250Å.

<sup>b</sup>Units of  $10^{-5}$

<sup>c</sup>Weighted mean

References. — Key to References – 1: Laurent, Vidal-Madjar, & York (1979); 2: Jenkins et al. (1999); 3: Hoopes et al. (2003); 4: Sonneborn et al. (2002); 5: Friedman et al. (2002)

In this paper, we report on our first set of Ti II observations of seven sightlines. We measure the Ti<sup>+</sup> column densities to assess the depletion levels along the sightlines and examine correlations with the observed D/H values. Finally, this Letter establishes a public database for Ti II measurements obtained by our group<sup>5</sup>. The data presented here, and all future observations, will be archived at this site, including the raw data and calibration frames. The data are freely available to other researchers.

## 2. OBSERVATIONS AND REDUCTION

The seven sightlines presented here were observed on the nights of September 8 and 9 and October 6, 2004 UT with the recently upgraded HIRES spectrometer (Vogt et al. 1994) on the 10m Keck I telescope. The spectrometer now contains a three CCD mosaic with  $> 90\%$  quantum efficiency at  $\lambda < 3500\text{\AA}$ . Because the targets are bright, we obtained the data primarily during twilight. The first half of the first night was marred by poor observing conditions while the remainder of the time was clear with typical seeing of  $\text{FWHM} \approx 0.7''$ . The targets in September 2004 were observed through the B5 decker ( $0.86''$  width;  $\text{FWHM} \approx 6\text{km s}^{-1}$  resolution;  $3.5''$  slit length) while BD+28°4211 was observed through the B1 decker ( $0.57''$  width;  $\text{FWHM} \approx 4.5\text{km s}^{-1}$  resolution;  $3.5''$  slit length). Table 1 summarizes the exposure times and data quality of the sample.

At the time of publication, a data reduction pipeline for the upgraded HIRES instrument was not available. Therefore, we reduced the two echelle orders containing the Ti II  $\lambda 3230, 3242, 3384$  transitions with a set of in-house IDL routines. These routines subtracted the bias, interactively set a boxcar aperture, interactively determined a region for scattered light subtraction (sky background was minimal),

<sup>5</sup> <http://www.ucolick.org/~xavier/TiII/index.html>

and extracted a 1D spectrum. A variance array was calculated accounting for the read noise and assuming Gaussian statistics. Wavelength calibration was carried out by fitting a 3rd order polynomial to the arc lines identified in the 1D spectra extracted from a ThAr image using the same trace and aperture as the science extraction. The typical RMS deviation from the fit was less than 0.1 pixels (i.e.  $< 0.003\text{\AA}$ ). Although we achieved a formal S/N ratio  $> 500 \text{pix}^{-1}$  for the brightest targets, the scatter around the stellar continuum significantly exceeds Gaussian statistics in the brightest sources. This is most likely due to noise in the combined flat field or perhaps small errors in the traces of the dispersed spectra. Throughout the analysis, we have augmented the variance arrays to match the empirical scatter measured in the stellar continuum.

All of the individual exposures were rebinned to a common wavelength scale with  $\lambda_0 = 3000\text{\AA}$  and a pixel size of  $1.4\text{km s}^{-1}$  after correcting to vacuum wavelengths and the heliocentric velocity reference frame. Multiple exposures of the same object were compared to identify cosmic rays and then coadded after scaling to a common flux and weighting by the square of the median signal-to-noise. Finally, we normalized the data by fitting a high order polynomial to the unfluxed stellar continuum. Continuum normalization is generally the largest source of uncertainty in the analysis, especially for the weakest transitions.

## 3. ANALYSIS

Figure 1 presents a velocity plot for the HD 191877 sightline. At the recorded S/N and resolution, it is straightforward to measure the Ti II ionic column densities. In the following, we assume that the majority of Ti<sup>+</sup> is in the  $^4\text{F}_{3/2}$  ground state. From our observations, we measure  $N(\text{Ti II } \lambda 3239) < 10^{10.6} \text{cm}^{-2}$  as a conservative upper limit to the  $^4\text{F}_{5/2}$  state ( $E \approx 135\text{K}$ ). Provided the

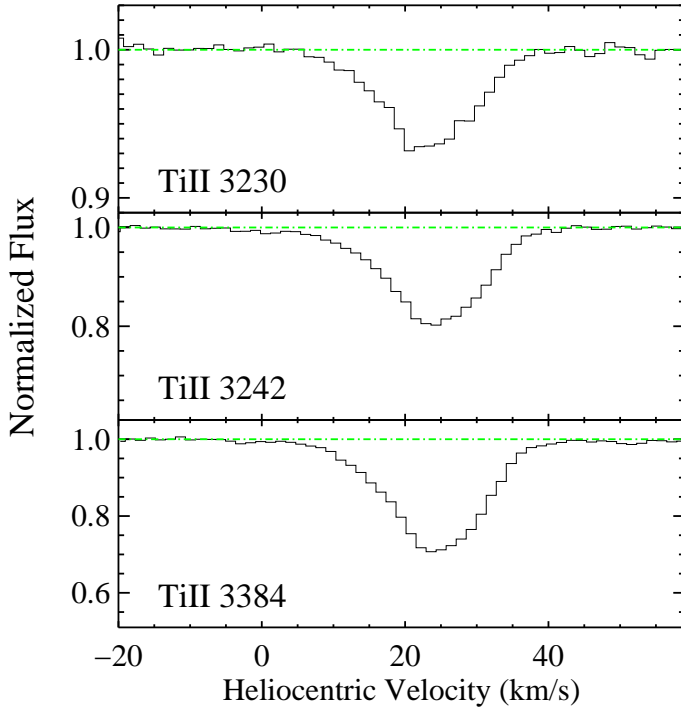


FIG. 1.— Ti II profiles for the ISM sightline toward HD 191877.

Ti II profiles have small optical depth, we can measure the Ti II column density by (i) summing the EW and assume the weak limit for the curve-of-growth (the values range from several to 30 mÅ); (ii) integrating the line profile with the apparent optical depth method (AODM; Savage and Sembach 1991; Jenkins 1996); and (iii) by fitting Voigt profiles to the data. All three techniques yield similar results; we have mainly used the AODM in this paper. The results for the three transitions<sup>6</sup> for all of the sightlines are listed in Table 1. For all measurements we report  $1\sigma$  uncertainties and the upper limits correspond to  $3\sigma$  limits. The uncertainties include (in quadrature) statistical uncertainty from Poisson noise and error due to continuum placement. We estimate a  $1\sigma$  continuum error of 0.1% for all of the objects except Feige 110 (0.2%) which corresponds to an approximately 0.5 mÅ error over the integrated profiles. Finally, we adopt a minimum error of 0.02 dex owing to systematic uncertainties related to data reduction (e.g. flat fielding).

Figure 2 presents  $\log(\text{Ti}/\text{H})$  against the  $(\text{D}/\text{H})$  values reported in the literature. The visual impression is suggestive of a correlation between the two quantities. The non-parametric Spearman and Kendall correlation tests reject the null hypothesis at the 96% and 95% c.l. respectively and indicate a positive correlation between the two quantities. This correlation provides preliminary evidence that the variation in Galactic D/H is related to the physical conditions that imply high depletion levels. Indeed, the results are consistent with the interpretation that low Galactic D/H values are due to an increased depletion of D onto grains.

Examining the figure, it is obvious that additional mea-

<sup>6</sup> Assuming  $f_{3230} = 0.0687$ ,  $f_{3242} = 0.232$ ,  $f_{3384} = 0.358$  (Bizzarri et al. 1993; Pickering, Thorne, & Perez 2001, 2002; Morton 2003).

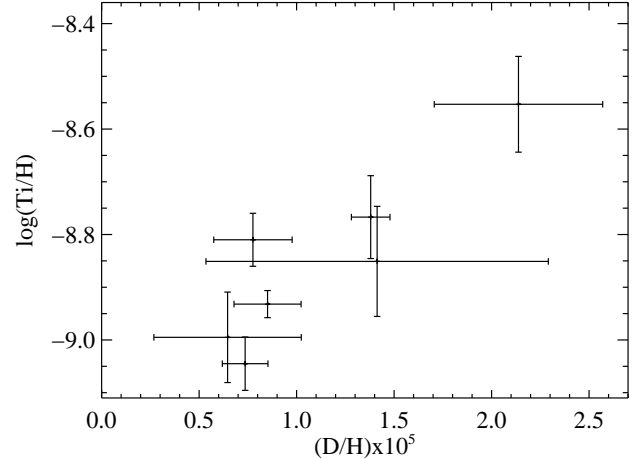


FIG. 2.—  $\log(\text{Ti}/\text{H})$  vs.  $\text{D}/\text{H}$  for the seven sightlines comprising the current sample. Error bars reflect  $1\sigma$  uncertainty.

surements of sightlines with large  $(\text{D}/\text{H})$  value will be especially valuable for testing the correlation. Furthermore, one observes a large spread in Ti/H at low D/H. If confirmed by future observations, the spread in Ti/H would argue against a simple linear correlation between Ti/H and D/H. We note that two additional stars ( $\zeta$  Pup,  $\gamma^2$  Vel) with  $(\text{D}/\text{H}) > 1 \times 10^{-5}$  have Ti II column densities reported in the literature (Welsh et al. 1997). The inferred Ti/H value for  $\zeta$  Pup ( $\log(\text{Ti}/\text{H}) = -8.5 \pm 0.08$ ) follows the trend indicated in Figure 2. The upper limit to Ti/H ( $< -8.9$ ) for  $\gamma^2$  Vel, however, suggests the gas along this sightline is significantly depleted even though its D/H value is among the highest known. A similar conclusion may be drawn from the GHRs observations of Fe II by Fitzpatrick & Spitzer (1994) although line-saturation is a potential concern. Although the confirmation of a high depletion level toward  $\gamma^2$  Vel would raise concern, Draine (2004a) notes that the processes of D depletion are different from those for other ions, e.g., D may be depleted in polycyclic aromatic hydrocarbons while heavier elements are depleted by other types of dust grains. Therefore one need not expect a one-to-one correspondence between Ti/H and D/H. In the coming year, we plan to acquire Ti II observations along many additional sightlines to better constrain the slope and scatter of Ti/H vs. D/H.

We have also examined the correlation between Ti/H and the  $N(\text{HI})$  values of the sightlines. Because one observes a correlation between volume density and depletion level (e.g. Jenkins 1987), one may expect a similar trend for  $N(\text{HI})$ . Furthermore, if photoionization is important along these sightlines its effects should correlate with  $N(\text{HI})$ . The Spearman and Kendall tests, however, reveal the null hypothesis is ruled out at only the 29% and 45% c.l. respectively. This follows the results from previous surveys of the Galactic disk Ti II (Stokes 1978; Welsh et al. 1997) although Wakker & Mathis (2000) see a strong correlation in a set of clouds with much larger dynamic range in  $N(\text{HI})$  than the sightlines considered here.

We wish to emphasize an important aspect of differential depletion in the analysis of Galactic D/H. In Figure 3, we present the Ti II profiles for the HD 195965 sightline

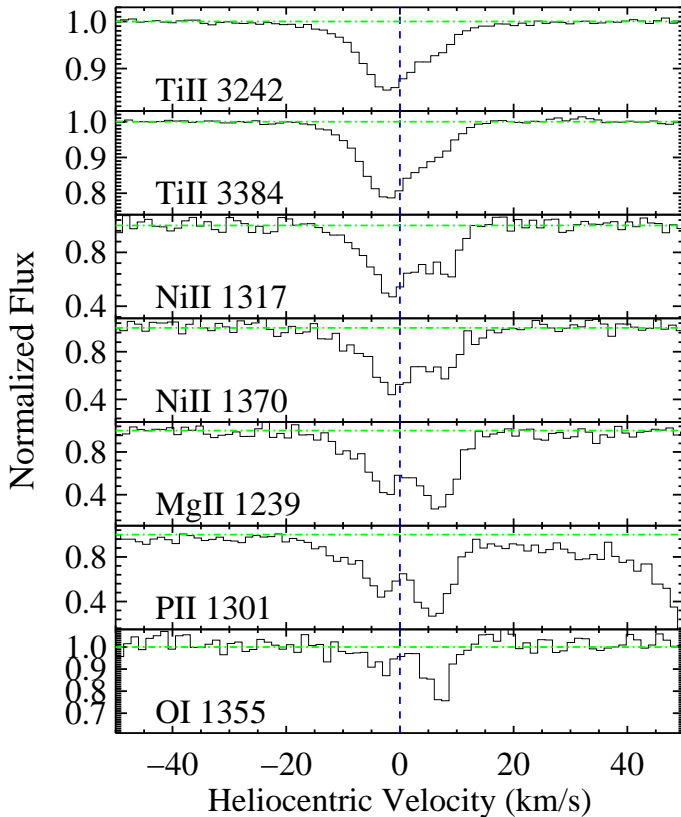


FIG. 3.— Velocity profiles for the ISM sightline toward HD 195965. Separately, the refractory ions ( $\text{Ni}^+$ ,  $\text{Ti}^+$ ) and non-refractory ions ( $\text{P}^+$ ,  $\text{O}^0$ ) show very similar sets of profiles. One notes, however, that the positive velocity component has a much larger depletion level for  $\text{Ni}^+$  and  $\text{Ti}^+$  ( $\approx 10\times$ ) than the negative component.

against a series of UV transitions obtained with STIS on the *Hubble Space Telescope*. The UV data have higher resolution ( $\text{FWHM} \approx 1.5 \text{ km s}^{-1}$ ) but lower S/N per resolution element than the optical observations. Examining the detailed component structure of the profiles, the refractory ions (e.g.  $\text{Ti}^+$ ,  $\text{Ni}^+$ ) have similar characteristics. Comparing against the non-refractory profiles ( $\text{P}^+$ ,  $\text{O}^0$ ), however, we identify the two main components but note that the relative abundances are significantly different. The Ti abundance relative to O is  $10\times$  greater in the positive component than the negative component. It is also noteworthy that the Mg II profiles more closely track the non-refractory species even though Mg is depleted along the sightline.

This raises a number of concerns regarding the analysis and interpretation of D/H. First, variations in the depletion level on a given sightline would generally lead to a weaker correlation between any existing Ti/H vs. D/H correlation if one only considers the integrated values. By a similar token, one would tend to underestimate the magnitude of intrinsic scatter in D/H regardless of the physical mechanism responsible. Another complication is that it is unclear whether one should constrain the analysis of D in lower resolution data (e.g. FUSE observations) with the velocity profiles traced by the O I, Mg II, or Ti II profiles. Profiles of Fe II are likely to be similar to Ni II and Ti II and could also be confusing in analysis of D I. A useful ex-

ercise would be to quantitatively compare high resolution D profiles with both refractory and non-refractory transitions along sightlines where the depletion levels vary. We consider these issues and examine the systematic effects on the inferred D/H ratios in a future paper. Finally, we note that the D profile of at least one extragalactic sightline shows significantly different component structure than the corresponding metal-line profiles (Tytler et al. 1996). Although the physical origin is more likely related to differences in metallicity or the ionization state of the metals, one cannot rule out intrinsic scatter in D/H even in these low metallicity sightlines. This issue could be relevant in studies of D/H in damped Ly $\alpha$  systems.

We acknowledge the efforts of S. Vogt, G. Hill, and the HIRES upgrade team for their efforts in completing the CCD mosaic upgrade. We would like to thank the referee B. Draine for helpful comments and criticism. Similarly, we appreciate the input from E. Jenkins, W. Moos, C. Oliveira, and K. Sembach. JXP is partially supported by NSF grant AST-0307408. TMT appreciates support from NASA through grant NNG04GG73G. This research was supported in part by the National Science Foundation under Grant No. PHY99-07949.

#### REFERENCES

- Bizzarri, A., Huber, M.C.E., Noels, A., Grevesse, N., Bereson, S.D., Tsekeris, P., & Lawler, J.E. 1993, *A&A*, 273, 707  
 Burles, S. and Tytler, D. 1998, *ApJ*, 507, 732  
 Chiappini, C., Renda, A., & Matteucci, F., 2002, *A&A*, 395, 789  
 Draine, B. 2004a, in “Origin and Evolution of the Elements”, ed. A. McWilliam and M. Rauch (astro-ph/0312592)  
 Draine, B. 2004b, FUSE Conference Proceeding, (astro-ph/0410310)  
 Fitzpatrick, E.L. & Spitzer, L.Jr. 1994, *ApJ*, 427, 232  
 Friedman, S.D., et al. 2002, *ApJS*, 140, 37  
 Hoopes, C.G., Sembach, K.R., Hébrard, G., Moos, H.W., & Knauth, D.C. 2003, *ApJ*, 586, 1094  
 Jenkins, E. B. 1996, *ApJ*, 471, 292  
 Jenkins, E.B. 1987, in *Interstellar Processes* ed. D.J. Hollenbach and H.A. Thronson, Jr. (Boston: D. Reidel Publishing Company), p. 533  
 Jenkins, E.B., Tripp, T.M., Wozniak, P.R., Sofia, U.J., & Sonneborn, G. 1999, *ApJ*, 520, 182  
 Jura, M. 1982, in *Advances in Ultraviolet Astronomy*, ed. Y. Kondo, NASA CP-2238, p. 54  
 Kingdon, J.B., & Ferland, G.J. 1996, *ApJS*, 106, 206  
 Kirkman, D., Tytler, D., Suzuki, N., O’Meara, J.M., & Lubin, D. 2003, *ApJS*, 149, 1  
 Laurent, C., Vidal-Madjar, A., & York, D.G. 1979, *ApJ*, 229, 923  
 Linsky, J.L., Diplas, A., Wood, B.E., Brown, A., Ayres, T.R., & Savage, B.D. 1995, *ApJ*, 451, L335  
 Lipman, K. & Pettini, M. 1995, *ApJ*, 442, 628  
 Morton, D.C. 2003, priv. comm.  
 O’Meara, J.M., Tytler, D., Kirkman, D., Nao, S., Prochaska, J.X., Lubin, D., & Wolfe, A.M. 2001, *ApJ*, 552, 718  
 Pickering, J.C., Thorne, A.P., & Perez, R. 2001, *ApJS*, 132, 403  
 Pickering, J.C., Thorne, A.P., & Perez, R. 2002, *ApJS*, 138, 247  
 Rogerson, J.B., Jr. & York, D.G.. 1976, *ApJ*, 186, L95  
 Savage, B. D. and Sembach, K. R. 1991, *ApJ*, 379, 245  
 Schramm, D.N. & Turner, S.M. 1998, *Rev. Mod. Phys.* 70., 303  
 Sonneborn, G. et al. 2002, *ApJS*, 140, 51  
 Spergel, D. et al. 2003, *ApJS*, 148, 175  
 Stokes, G.M. 1978, *ApJS*, 36, 115  
 Tytler, D., Fan, X.-M., & Burles, S. 1996, *Nature*, 381, 207  
 Vogt, S.S., Allen, S.L., Bigelow, B.C., Bresee, L., Brown, B., et al. 1994, *SPIE*, 2198, 362  
 Wakker, B.P. & Mathis, J.S. 2000, *ApJ*, 544, L107  
 Welsh, B.Y., Sasseen, T., Craig, N., Jelinsky, S., & Albert, C.E. 1997, *ApJS*, 112, 507  
 Wood, B.E., Linsky, J.L., Hébrard, G., Williger, G.M., Moos, H.W., & Blair, W.P. 2004, *ApJ*, 609, 838